

INFERRING INITIAL SPIN PERIODS FOR NEUTRON STARS IN COMPOSITE REMNANTS

E. VAN DER SWALUW¹, Y. WU²*Draft version June 4, 2001*

ABSTRACT

We propose a method to infer the initial spin period of pulsars residing in composite supernova remnants. Such a remnant consists of both a plerionic and a shell type component, corresponding respectively to the pulsar wind nebula driven by the spindown luminosity of the central pulsar, and the blastwave bounding the supernova remnant. Theoretical investigations including hydrodynamical simulations have shown that at late times ($\sim 1,000$ - $10,000$ years), a simple scaling law connects the radius of the supernova shell to the radius of the plerion. The energy content of the plerion and the total mechanical energy of the supernova remnant enter into this scaling law. One can use this scaling law to estimate the initial spin period of pulsars residing in composite remnants. We discuss potential pitfalls of this method, including the effect of a small remnant age and of strong radiative losses in the plerion.

Subject headings: hydrodynamics—(ISM:) supernova remnants—(stars:) pulsars: general

1. INTRODUCTION

Is a typical neutron star born spinning close to break-up ($P_0 \sim 0.8$ ms, Haensel et al. 1995). The answer to this question reveals much about the angular momentum evolution of neutron star progenitors (Heger et al. 2000), as well as about the processes that occur during the supernova explosion (Spruit & Phinney 1998, Lai & Goldreich 2000). If the initial spin-rate is close to the break-up rate, the nascent neutron star may undergo a Rossby-wave instability (Lindblom, Owen & Morsink 1998) and emit gravitational radiation. Moreover, a rapidly rotating neutron star can obtain a substantial kick velocity through the electromagnetic rocket effect (Harrison & Tademaru 1975). This can provide a natural explanation for the apparent alignment between the spin axis and pulsar proper motion (Lai et al. 2001) in the Crab (Caraveo & Mignani 1999) and Vela pulsar (Pavlov et al. 2000).

A number of studies have been dedicated to the determination of initial spin-rates. Those few young pulsars which have measured braking indices and estimated ages seem to possess initial spin periods from 19 ms (Crab, Lyne et al. 1993) to 63 ms (PSR B1509-58, Kaspi et al. 1994).³ Since only rapidly rotating pulsars have measurable braking indices, these values may not be representative for the general population. For these, Phinney & Blandford (1981) and Vivekanand & Narayan (1981) have developed the pulsar-current analysis. Unfortunately, this method is subject to small-number statistics (Lorimer et al. 1993). Analysis based on pulsar luminosity functions (Emmering & Chevalier 1989, Narayan 1987) raised the possibility of ‘injection’: the majority of the pulsars may be born with periods as slow as several hundred milliseconds. This hypothesis is corroborated by the apparent paucity of plerions around neutron stars (Srinivasan et al. 1984). A plerion derives energy from the rotational energy of the central pulsar. The overall conclusion from pulsar studies seems to favour a slow initial spin rate.

In this article, we propose a new method for inferring the initial spin-rates of some neutron stars – the small population that reside within composite supernova remnants. A compos-

ite supernova remnant (SNR) includes both a plerionic (filled-centered) component (Weiler & Panagia 1978) and a shell component. The first corresponds to the pulsar wind nebula (PWN), while the latter corresponds to the blastwave of the supernova remnant (SNR) propagating through the interstellar medium (ISM). They are observationally distinguished by a different spectral power-law index at radio frequencies. The dynamics of such a composite system has been extensively studied, e.g. Reynolds & Chevalier (1984); Chevalier & Fransson (1992).

Recently, van der Swaluw et al. (2001) presented hydrodynamical simulations and analytical arguments which explicitly relate the radius of the PWN to that of the SNR. These authors find that the two radii are roughly proportional some time after the initial explosion, with the proportionality constant determined by the ratio of the pulsar spin-down energy and the total mechanical energy of the supernova event. This relation forms the theoretical basis for the current paper. In this article, we follow van der Swaluw et al. (2001) in restricting ourselves to spherically symmetric systems.

2. EVOLUTION OF A PWN IN A SNR

As a pulsar spins down, its rotational energy is largely deposited into the surrounding medium, driving a relativistic pulsar wind into the tenuous bubble of stellar debris left behind by the supernova blastwave. We approximate the spin-down luminosity of the pulsar wind by that of a rotating magnetic dipole:

$$L(t) = L_0 / (1 + t/\tau)^2. \quad (1)$$

The total spin-down energy is $E_{sd} = L_0 \tau$. We neglect the dynamical influence of the pulsar wind on the SNR, assuming that E_{sd} is well below the total mechanical energy of the supernova explosion, $E_0 \sim 10^{51}$ erg.

The evolution of a PWN inside a SNR can be divided into two stages (Reynolds & Chevalier 1984; Chevalier & Fransson 1992). In the first stage, the PWN expands supersonically within the bubble blown out by the SNR. The second stage commences at $t \equiv t_{ST}$ when the PWN encounters the reverse shock of the SNR. The reverse shock heats the interior of the

¹E.vanderswaluw@astro.uu.nl, Astronomical Institute, Utrecht University, P.O. Box 80000, 3508 TA Utrecht, the Netherlands

²wu@cita.utoronto.ca, Canadian Institute for Theoretical Astrophysics, University of Toronto, 60 St. George Street, Toronto, Ontario M5S 3H8, Canada

³The LMC x-ray pulsar J0537-6910 (16 ms, Marshall et al. 1998) should be born at shorter than 10 ms for reasonable values of the braking index.

SNR and causes the expansion of the PWN to become subsonic. In this stage the expansion of the PWN is regulated by the expansion of the SNR, which is described by a self-similar Sedov-Taylor solution. In the following we summarize some relevant results from Chevalier & Fransson (1992) and van der Swaluw et al. (2001).

2.1. Analytical Relations

We focus on the subsonic expansion stage of the PWN which occurs when the SNR has relaxed to the Sedov solution. We assume that the pulsar wind has deposited most of its energy into the PWN ($t \gg \tau$). We find that the radius of the PWN (R_{pwn}) scales roughly linearly with the radius of the SNR (R_{snr}). This scaling arises from the condition of pressure equilibrium between the PWN interior and the interior of the SNR. We approximate the pressure which confines the PWN by the central pressure of the SNR, using the Sedov solution (Sedov 1958),

$$P_{\text{snr}} \simeq 0.074 E_0 / R_{\text{snr}}^3, \quad (2)$$

where E_0 is the total mechanical energy of the SNR. Inside the PWN, pressure is quickly equilibrated due to the high sound speed $\sim c/\sqrt{3}$ of the relativistic fluid. The interior pressure is related to the energy content of the PWN, E_* , by

$$P_{\text{pwn}} \simeq \frac{3(\gamma-1)}{4\pi} \frac{E_*}{R_{\text{pwn}}^3}, \quad (3)$$

where γ is the adiabatic index. Let E_{pwn} be the total amount of energy injected into the PWN by the central pulsar. Part of this energy is used to perform work on the surrounding medium as the PWN expands. Despite some compression of the PWN by the reverse shock around $t = t_{\text{ST}}$, E_* falls below E_{pwn} (see Fig. 9 of van der Swaluw et al. (2001)). We therefore write $E_* = \eta_2 E_{\text{pwn}}$ with $\eta_2 < 1$ decreasing over time. Imposing pressure equilibrium, we find

$$R_{\text{pwn}} = \bar{C} (E_*/E_0)^{1/3} R_{\text{snr}}, \quad (4)$$

where $\bar{C} \simeq 1.02$ for a relativistic fluid ($\gamma = 4/3$) and $\bar{C} \simeq 1.29$ for a non-relativistic fluid ($\gamma = 5/3$).

In the absence of radiative losses, the subsonic expansion of the PWN is adiabatic, $P_{\text{pwn}} V_{\text{pwn}}^\gamma = \text{constant}$, where $V_{\text{pwn}} \propto R_{\text{pwn}}^3$ is the volume of the PWN. Combining this relation with equation (3), we find that E_* falls off with time as $t^{-3/10}$ for a relativistic fluid. This yields an expansion law for the radius of the PWN, $R_{\text{pwn}} \propto t^{3/10}$, whereas the radius of the SNR scales as in a Sedov solution ($R_{\text{snr}} \propto t^{2/5}$). So, roughly, $R_{\text{pwn}} \propto R_{\text{snr}}$.

We introduce two additional dimensionless parameters, η_1 and η_3 , and rewrite equation (4) into a form that can be more easily compared with simulations,

$$R_{\text{pwn}}(t) = \eta_3(t) (\eta_1 E_{\text{sd}}/E_0)^{1/3} R_{\text{snr}}(t). \quad (5)$$

The parameter η_1 is a constant and relates the total energy input into the PWN, E_{pwn} , to the total spin-down energy of the pulsar, E_{sd} as $E_{\text{pwn}} = \eta_1 E_{\text{sd}}$. This takes into account possible radiative losses in the PWN, and other inefficiencies in the conversion of the spin-down energy into mechanical energy. For instance, a fraction of the neutron star spin-down energy may escape directly from the pulsar as high energy radiation. We set $\eta_1 = 1$ in this article, but discuss the case when it is much less than

unity. The second parameter is defined as $\eta_3(t) = \bar{C} \eta_2^{1/3} \leq \bar{C}$. The maximum value of $\eta_3 = \bar{C}$ is used when we determine the initial spin periods of pulsars driving composite remnants.

The derivation leading up to equation (5) is invalid during the earlier supersonic evolution. We retain equation (5) with the sole purpose of defining η_3 during these stages. We proceed to demonstrate the behaviour of η_3 throughout the PWN evolution with results from hydrodynamical simulations.

2.2. Hydrodynamical Simulations

Van der Swaluw et al. (2001) have presented hydrodynamical simulations of a pulsar wind nebula inside a supernova remnant. In these simulations, the gasdynamical equations are integrated in a spherically symmetric configuration, using the Versatile Advection Code (VAC) developed by Gábor Tóth at the Astronomical Institute Utrecht (Tóth and Odstrčil 1996). Lacking the possibility of treating a relativistic fluid, the pulsar wind has been implemented as a cold non-relativistic wind ($\gamma = 5/3$) with a terminal velocity equal to the speed of light, $v_\infty = [2L(t)/\dot{M}_{\text{pw}}(t)]^{1/2} \simeq c$, where $\dot{M}_{\text{pw}}(t)$ is the mass ejection rate into the pulsar wind bubble, and $L(t)$ is the pulsar spin-down luminosity as given by equation (1).

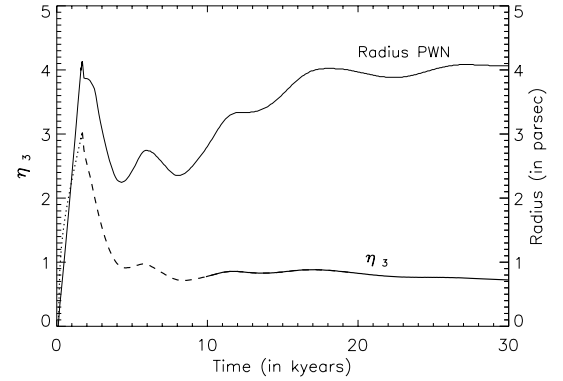


FIG. 1.— Results of numerical simulation showing the time evolution of the pulsar wind nebula. Plotted here are its radius (in units of parsec, thin continued line) and the dimensionless parameter η_3 (thick line). The latter is dissected into three parts: a supersonic expansion stage when the PWN is bounded by a shock (dotted section), in unsteady transition when undergoing reverberation with the supernova reverse shock (dashed section), and subsonic expansion (solid section). The parameters adopted in this simulation is: maximum spin-down luminosity $L_0 = 5 \times 10^{38}$ ergs/s, spin-down time $\tau = 600$ yrs, supernova explosion energy of $E_0 = 10^{51}$ ergs and ISM density of 10^{-24} g/cm³. We find $t_{\text{ST}} \sim 10^3$ yrs in this case. Note that η_3 is not physical in the early expansion stages but is only formally defined by equation (5). For details, see van der Swaluw et al. (2001).

Figure 1 presents relevant results from one of these simulations. The numerical parameters adopted in this simulation are listed in the caption. Taking different parameters will not qualitatively change the overall behaviour of the system, but it will affect the early (supersonic) evolution and the moment when subsonic expansion commences (t_{ST}). For example, t_{ST} is larger if the supernova explosion is more energetic, if the supernova ejecta mass is larger, or if the interstellar medium density is smaller ($t_{\text{ST}} \propto E_0^{-1/2} M_{\text{ej}}^{5/6} \rho_{\text{ism}}^{-1/3}$, e.g. McKee & Truelove (1995)). Since this is a non-relativistic simulation ($\gamma = 5/3$), we expect $\eta_3 \leq \bar{C} = 1.29$ in the subsonic expansion stage. This is indeed observed. Therefore in the following investigation we assume a maximum value for η_3 of 1.02, as is appropriate for the subsonic expansion stage of a PWN in a SNR with $\gamma = 4/3$.

3. INFERRING INITIAL SPIN RATES

Writing the spin-down energy as $E_{\text{sd}} = (\Omega_0^2 - \Omega_t^2)I/2$, we obtain the following expression for the pulsar's initial spin period,

$$P_0 = 2\pi \left(\frac{2E_0}{\eta_1 I} \left(\frac{R_{\text{pwn}}}{\eta_3 R_{\text{snr}}} \right)^3 + \left(\frac{2\pi}{P_t} \right)^2 \right)^{-1/2}, \quad (6)$$

where the spin-periods are $P_0 \equiv 2\pi/\Omega_0$, $P_t \equiv 2\pi/\Omega_t$ and I is the moment of inertia of the neutron star. The radius of a PWN relative to its associated SNR shell can be used to infer the initial spin-period of the pulsar, assuming η_1 , η_3 , E_0 and Ω_t are known or can be estimated.

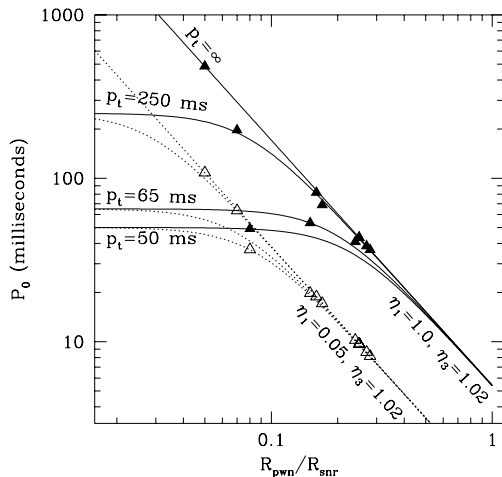


FIG. 2.— Initial spin periods inferred from the ratio $R_{\text{pwn}}/R_{\text{snr}}$ using equation (6). When assuming no radiative loss ($\eta_1 = 1$), we obtain solutions as depicted by the family of solid lines, with different lines applicable for different current spin periods, P_t . The family of dashed lines are for $\eta_1 = 0.05$. Filled and open triangles represent systems listed in Table 1 for $\eta_1 = 1$ and $\eta_1 = 0.05$, respectively. When $P_0 \leq 4$ ms, the initial rotational energy of the pulsar exceeds the total mechanical energy in the SNR itself. Equation (6) fails and the SNR may be blown away by the PWN. None of the systems we examined lies close to this limit.

The radius of PWN should be taken from radio measurements. Synchrotron-radiating electrons in PWN have a typical life-time (Rybicki & Lightman 1979):

$$\tau_{1/2} = \frac{6\pi m_e c}{\sigma_T B^2 \Gamma} \sim 10^5 B_{\text{mG}}^{-3/2} \nu_{\text{MHz}}^{-1/2} \text{ yrs}, \quad (7)$$

where σ_T is the Thomson cross-section, Γ is the electron Lorentz factor, ν is the typical synchrotron frequency scaled at the radio wavelength, and B is the magnetic field strength scaled by a value appropriate for plerions. Such a long life-time implies that radio observations ‘see’ electrons that were accelerated during the early energetic stage when the pulsar deposits most of its rotational energy into the pulsar wind. They therefore delineates the true spatial extent of the plerion. In contrast, X-ray electrons lose their energy quickly. The X-ray plerion is produced by electrons recently accelerated near the pulsar, so one expects the plerion to be smaller in X-rays.

We have collected from the literature 13 composite supernova remnants. We list the relevant properties of these systems in Table 1. The relative sizes of the plerions are taken from radio observations, with the exception of G11.2–0.3 and G320.4–1.2 (marked by asterisks), where only X-ray plerions have been observed. We apply equation (6) to determine the initial spin period for these systems, adopting $\eta_1 = 1.0$, $\eta_3 = 1.02$,

$E_0 = 1.0 \times 10^{51}$ erg, $I = 1.4 \times 10^{45}$ g cm² (the value for Crab), and $P_t = 2\pi/\Omega_t = \infty$ in cases where it is unknown. The same results appear in Figure 2 as filled triangles. By taking the maximum value for η_3 we probably under-estimate the initial spin rate, but likely by no more than a factor of 2 judging from Figure 1. In the same table we also list the initial spin periods (P'_0) derived for 4 pulsars using braking index measurements.

4. DISCUSSION

We have demonstrated a new method for inferring initial spin rates of pulsars residing in composite supernova remnants. This method uses the ratio of the plerion radius and the radius of the supernova shell, so we dub it as the ‘radius method’. Given that typical plerions and supernova remnants are expected to live for 10^3 – 10^4 years, we expect that future observations will increase the population to which the radius method can be applied.

This method does not require a knowledge of parameters like the density and magnetic field strength of the ISM, or the distance and age of the SNR. However, for the method to apply, the observed PWN has to be in the subsonic stage of its evolution. If this is not the case, the method systematically over-estimates the initial spin rate (see below). Uncertainties in E_0 , η_3 and Ω_t introduce errors of order unity, as will a deviation from pure spherical symmetry. A more important uncertainty of this new method concerns the influence of synchrotron radiative losses in the plerion, as parametrized by η_1 . We tend to under-estimate the initial spin-rate by ignoring these losses (setting $\eta_1 = 1$). We will illustrate the effect of these uncertainties by considering the 4 systems in Table 1, which have known central pulsars, and independent estimates for their initial periods based on measurements of the pulsar braking index.⁴

Among these systems, Vela is the oldest and should have entered the subsonic expansion stage. Its associated pulsar being displaced from the center of the plerion also testifies in favor of an old age. Adopting $\eta_3 = 1.02$ and other parameters as stated, the radius method yields an initial spin-down energy for Vela that is a factor of 2 above that inferred by Lyne et al. (1996) from pulsar braking index measurement. This is surprising as we have ignored radiative loss. We suggest three possibilities that will reconcile the two methods: (1) the supernova explosion is half as energetic as assumed here, (2) the pulsar is 15 kyr old, and (3) the braking index is as high as 2.4 (Lyne et al. 1996 reported 1.4). The first option will raise P_0 to 52 ms while the latter two will reduce P'_0 to 41 ms.

The remaining two systems, G11.2–0.3 and G320.4–1.2, are young (~ 1.5 kyr). Their plerions could still be in supersonic expansion during which the value of η_3 (defined formally through equation (5) in the supersonic stage) can be higher than 1.02, the theoretical maximum during the subsonic expansion stage (Fig. 1 and §2 of van der Swaluw et al. 2001). For example, taking $\eta_3 = 2$ we find $P_0 = 63$ ms for G11.2–0.3. A similar correction may have to be applied to G29.7–0.3 (Kes 75).

In contrast to the systems discussed above, SNR 0540–69.3 is the only case in which the spin-down energy is significantly under-estimated for our choice of parameters. The plerion in SNR 0540–69.3 is too small if the pulsar was indeed born with an initial spin of 39 ms. Manchester et al. (1993) have argued that the radiative loss in this system has reduced the total energy stored in the plerionic magnetic field and relativistic electrons to a mere 4% of the integrated spin-down luminosity of the pulsar. Taking $\eta_1 = 0.04$ we recover a 39 ms initial period, with a

⁴G11.2–0.3 does not have a braking index measurement. But its inferred initial spin depends little on this index due to its small age.

Supernova Designation	Plerion $R_{\text{pwn}}/R_{\text{snr}}$	P_0 (in msec)	Pulsar P_t (in msec)	P'_0 (in msec)	age (kyr)
G0.9+0.1	0.25(Helfand & Becker 1987)	43	—	—	—
G16.7+0.1	0.25(Helfand et al. 1989)	43	—	—	—
G29.7−0.3	0.16(Becker & Helfand 1984)	82	325[PSR 1846-0258]	—	0.7
G34.7−0.4	0.07(Frail et al. 1996)	197	267[PSR 1853+01]	—	—
G293.8+0.6	0.25(Whiteoak & Green 1996)	43	—	—	—
G322.5−0.1	0.27(Whiteoak 1992)	39	—	—	—
G326.3−1.8	0.28(Kassim et al. 1993)	37	—	—	—
G327.1−1.1	0.25(Sun et al. 1999)	43	—	—	—
G351.2+0.1	0.05(Becker & Helfand 1988)	484	—	—	—
G263.9−3.3	0.24(Milne 1980)	41	89[Vela]	52(Lyne et al. 1996)	11.3
G11.2−0.3*	0.15(Vasisht et al. 1996)	53	65	63(Torii et al. 1999)	1.6
G320.4−1.2*	0.17(Seward et al. 1983)	69	150[PSR B1509-58]	63(Kaspi et al. 1994)	1.5
SNR0540−69.3	0.08(Manchester et al. 1993)	49	50.3[PSR 0540-69]	39(Deeter et al. 1999)	0.8

TABLE 1

LIST OF COMPOSITE REMNANTS: P_0 IS THE INITIAL SPIN PERIOD FROM THE RADIUS METHOD, P_t IS THE CURRENT SPIN PERIOD OF SYSTEMS WITH IDENTIFIED PULSARSQ AND P'_0 IS THE CALCULATED INITIAL SPIN PERIOD BY USING THE BRAKING INDEX.

value of $\eta_3 = 1.3$. This value of η_3 is reasonable because SNR 0540−69.3 is only 760 years old. None of the other 3 systems discussed here requires strong synchrotron losses in the plerion, which makes SNR 0540−69.3 unique. Among the 4 pulsars, PSR 0540−69 is not outstanding either in its surface magnetic field strength or in its spin-down rate. It is likely born the fastest, though not by a large margin.

Along the same line, Atoyan (1999) have argued that in Crab, strong synchrotron losses in its past can account for the observed flat spectra of the radio electrons in the plerion. If this applies to all PWNe, it would significantly change our inferred initial spin rates. This is illustrated in Fig. 2 for a strong synchrotron loss of $\eta_1 = 0.05$.

Although the effect of radiative losses on the reliability of our results remain unclear, we briefly explore the implications of these results. Our sample (table 1) is likely biased towards fast

spinners, which produce plerions that are easier to detect. Even so, all pulsars in our sample seem to be born spinning much below the break-up rate. Moreover, their initial spin-periods do not cluster around $\sim 5 - 10$ ms, a signature of rapid spin-down by the gravitationally radiating Rossby-waves which occur immediately after the supernova collapse (Andersson et al. 2000, Wu et al. 2001). The long initial periods we find imply an effective angular momentum coupling between the core and envelope of neutron star progenitors until late into their evolution. The scattering in the initial periods could either reflect a different decoupling time in different stars, or a stochastic process that gives rise to the final angular momentum in the neutron star, such as an off-centered kick during the collapse (Spruit & Phinney 1998).

We acknowledge discussions with Bram Achterberg, Armen Atoyan and an anonymous referee.

REFERENCES

- Andersson, N., Jones, D. I., Kokkotas, K. D. & Stergioulas, N. 2000, *ApJ*, 534, L75
- Atoyan, A. M. 1999, *A&A*, 346, L49
- Becker, R. H. & Helfand, D. J. 1984, *ApJ*, 283, 154
- Becker, R. H. & Helfand, D. J. 1988, *AJ*, 95, 883
- Caraveo, P. A. & Mignani, R. P. 1999, *A&A*, 344, 367
- Chevalier, R.A., Fransson, C., 1992, *ApJ*, 395, 540
- Deeter, J. E., Nagase, F. & Boynton, P. E. 1999, *ApJ*, 512, 300
- Emmering, R. T. & Chevalier, R. A. 1989, *ApJ*, 345, 931
- Frail, D. A., Giacani, E. B., Goss, W. M. & Dubner, G. 1996, *ApJ*, 464, L165
- Haensel, P., Salgado, M., & Bonazzola, S. 1995, *A&A*, 296, 745
- Harrison, E. R. & Tademaru, E. 1975, *ApJ*, 201, 447
- Heger, A., Langer, N. & Woosley, S. E. 2000, *ApJ*, 528, 368
- Helfand, D. J. & Becker, R. H. 1987, *ApJ*, 314, 203
- Helfand, D. J., Velusamy, T., Becker, R. H. & Lockman, F. J. 1989, *ApJ*, 341, 151
- Kassim, N. E., Hertz, P. & Weiler, K. W. 1993, *ApJ*, 419, 733
- Kaspi, V. M., Manchester, R. N., Siegman, B., Johnston, S. & Lyne, A. G. 1994, *ApJ*, 422, L83
- Lai, D. & Goldreich, P. 2000, *ApJ*, 535, 402
- Lai, D., Chernoff, D. F. & Cordes, J. M. 2001, *ApJ*, 549, 1111
- Lindblom, L., Owen, B. J. & Morsink, S. M. 1998, *Phy. Rev. L.* 80, 4843
- Lorimer, D. R., Bailes, M., Dewey, R. J. & Harrison, P. A. 1993, *MNRAS*, 263, 403
- Lyne, A. G., Pritchard, R. S. & Graham-Smith, F. 1993, *MNRAS*, 265, 1003
- Lyne, A. G., Pritchard, R. S., Graham-Smith, F. & Camilo, F. 1996, *Nature*, 381, 497
- Manchester, R. N., Staveley-Smith, L. & Kesteven, M. J. 1993, *ApJ*, 411, 756
- Marshall, F. E., Gotthelf, E. V., Zhang, W., Middleditch, J. & Wang, Q. D. 1998, *ApJ*, 499, L179
- McKee, C.F. & Truelove, Phys.Rep. 256,157
- Milne, D. K. 1980, *A&A*, 81, 293
- Narayan, R. 1987, *ApJ*, 319, 162
- Pavlov, G. G., Sanwal, D., Garmire, G. P., Zavlin, V. E., Burwitz, V. & Dodson, R. G. 2000, AAS meeting, 196, 3704
- Phinney, E. S. & Blandford, R. D. 1981, *MNRAS*, 194, 137
- Rybicki, G. B., Lightman, A. P. 1979, *Radiative Processes in Astrophysics*, John Wiley & Sons,
- Reynolds, S.P., and Chevalier, R.A., 1984, *ApJ*, 278, 630
- Sedov, L.I., 1959, *Similarity and Dimensional Methods in Mechanics*, Academic Press, New York
- Seward, F. D., Harnden, F. R., Murdin, P. & Clark, D. H. 1983, *ApJ*, 267, 69
- Seward, F. D. & Wang, Z. R. 1988, *ApJ*, 332, 199
- Spruit, H. C. & Phinney, E. S. 1998, *Nature*, 393, 139
- Srinivasan, G., Dwarkanath, K. S. & Bhattacharya, D. 1984, *J. of Astrophysics. and Astronomy*, 5, 403
- Sun, M., Wang, Z. & Chen, Y. 1999, *ApJ*, 511, 274
- Tóth, G. & Odstrčil, 1996 *J. Comp. Phys.* 128, 82
- Torii, K., Tsunemi, H., Dotani, T., Mitsuda, K., Kawai, N., Kinugasa, K., Saito, Y. & Shibata, S. 1999, *ApJ*, 523, L69
- van der Swaluw, E., Achterberg, A., Gallant, Y. A. & Tóth, G. 2001, *astro-ph/0012440*, submitted to *A&A*
- Vasisht, G., Aoki, T., Dotani, T., Kulkarni, S. R. & Nagase, F. 1996, *ApJ*, 456, L59
- Vivekanand, M. & Narayan, R. 1981, *J. of Astrophysics and Astronomy*, 2, 315
- Weiler K. W. & Panagia N., 1978, *A&A* 70, 419
- Whiteoak, J. B. Z. 1992, *MNRAS*, 256, 121
- Whiteoak, J. B. Z. & Green, A. J. 1996, *A&AS*, 118, 329
- Wu, Y., Matzner, C. D., Arras, P. 2001, *ApJ*, 549, 1011